

MWA array simulation with MAPS

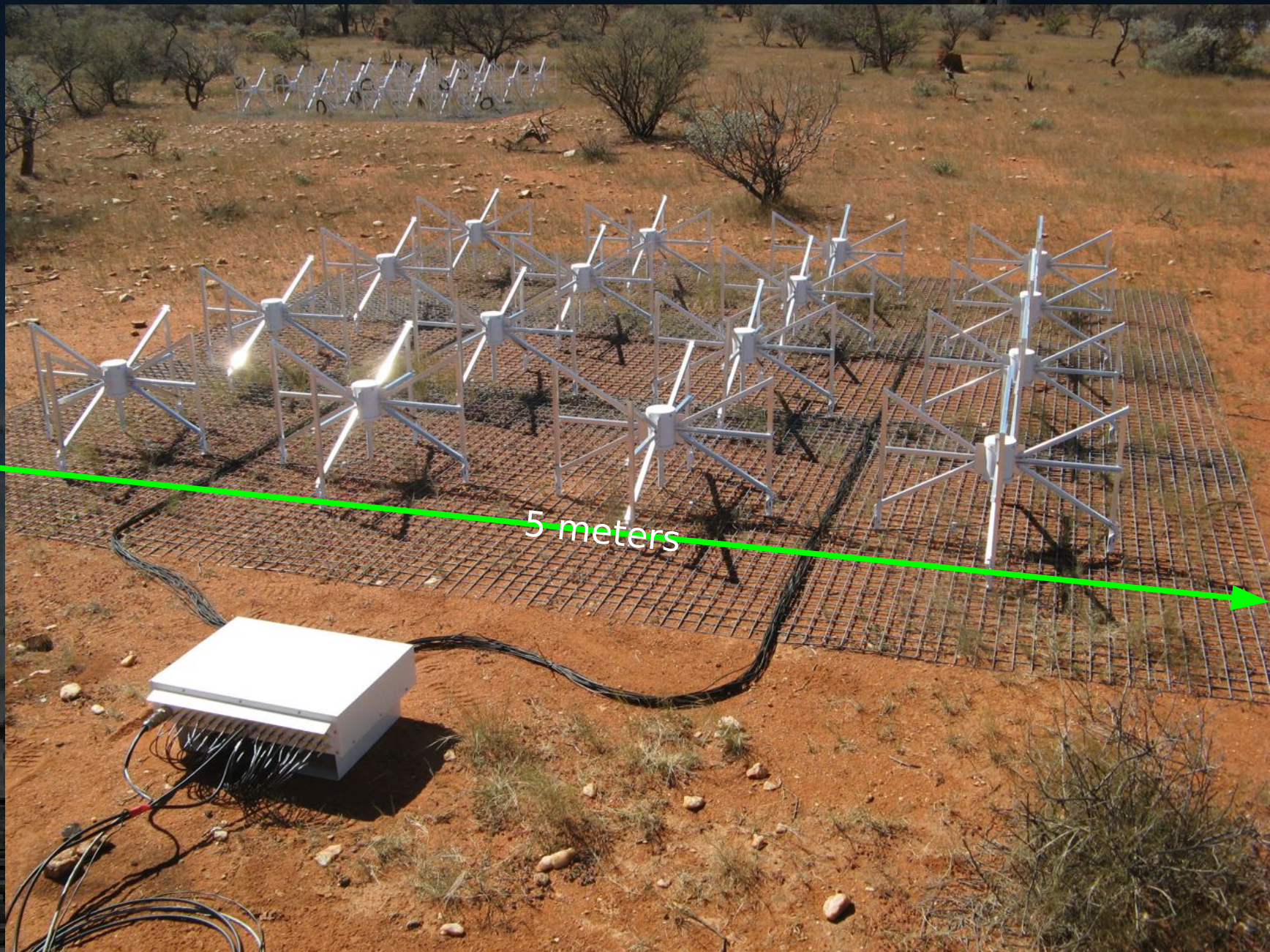
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Harvard-Smithsonian Center for Astrophysics

Array Specs

- 512 antennas
 - 15-50° field of view
- A wide frequency range
 - 80 - 300 MHz, 31MHz at any one time
- Full cross correlation of all 512 antennas
 - full Stokes, 31 MHz bandwidth, 40 kHz resolution

MWA antennas



Coplanar enough for you?






Photo: Divya Oberoi

Simulations - 'MAPS'

- MAPS = MIT Array Performance Simulator
- MAPS simulates what an interferometer sees, from ionosphere through to correlator including station beams
- Original development at MIT Haystack (Cappallo, Doeleman, Lonsdale, Oberoi, et al.)
- Significant enhancements for MWA (Wayth, Mitchell, Ord, Kasper):
 - antenna beams
 - polarized response
 - all-sky

Simulator Status

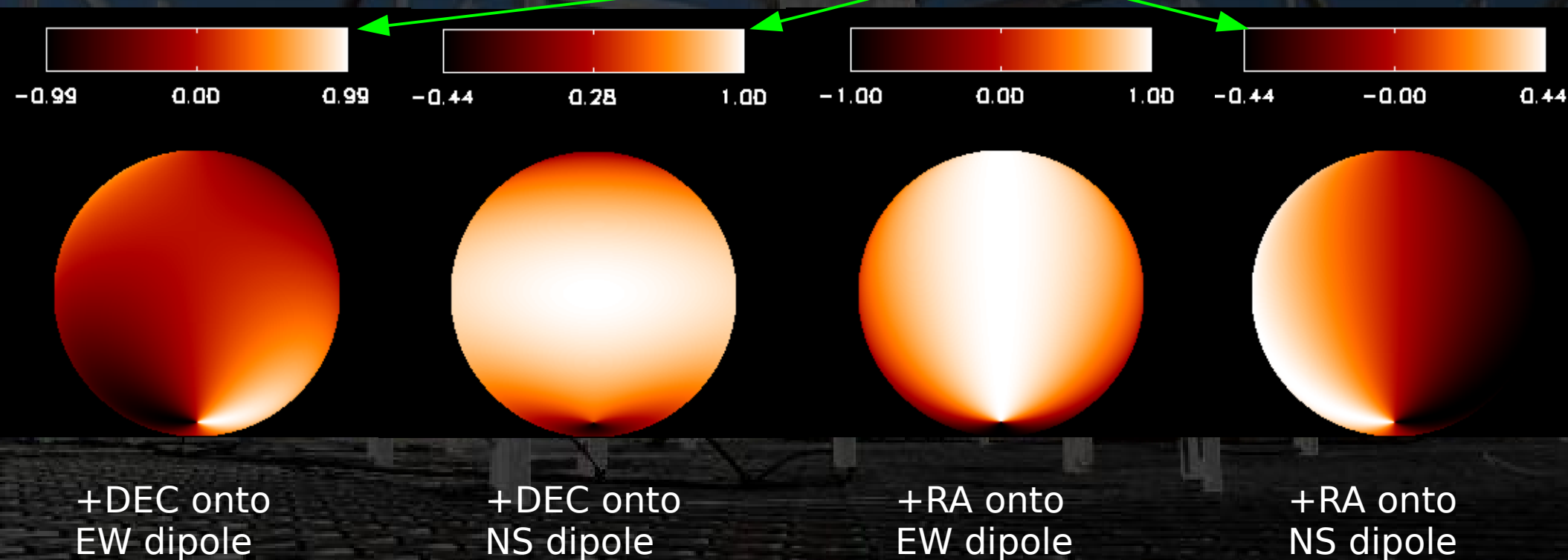
- Fully polarized instrument response
- Realistic, non-equal, dual-pol antennas
- Sky comprising diffuse and point sources with polarized description of sky 
- All sky coverage
- Different pointing and phase centers 
- Ionosphere model with large scale density and small scale turbulent structure. 
- Ionospheric Faraday rotation



Fixed dipole polarised response

- For lat -26°
- dipoles aligned with E-W or N-S
- Response to E field aligned with RA or DEC

Note **linear** scales



+DEC onto
EW dipole

+DEC onto
NS dipole

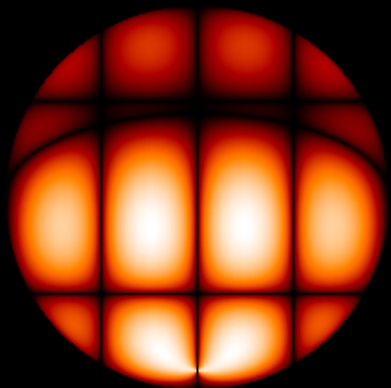
+RA onto
EW dipole

+RA onto
NS dipole

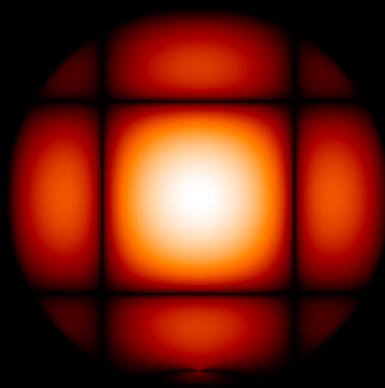
MWA antenna Jones matrices

- Phased arrays overlay array pattern and ground screen attenuation.
- Examples at 140MHz

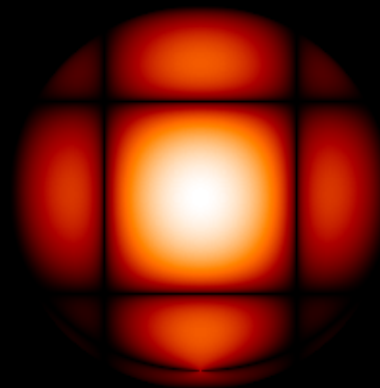
Note dB scales!



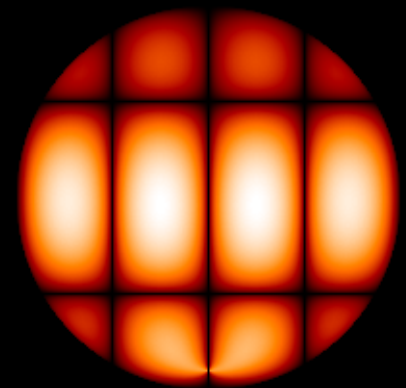
+DEC onto
EW antenna



+DEC onto
NS antenna



+RA onto
EW antenna

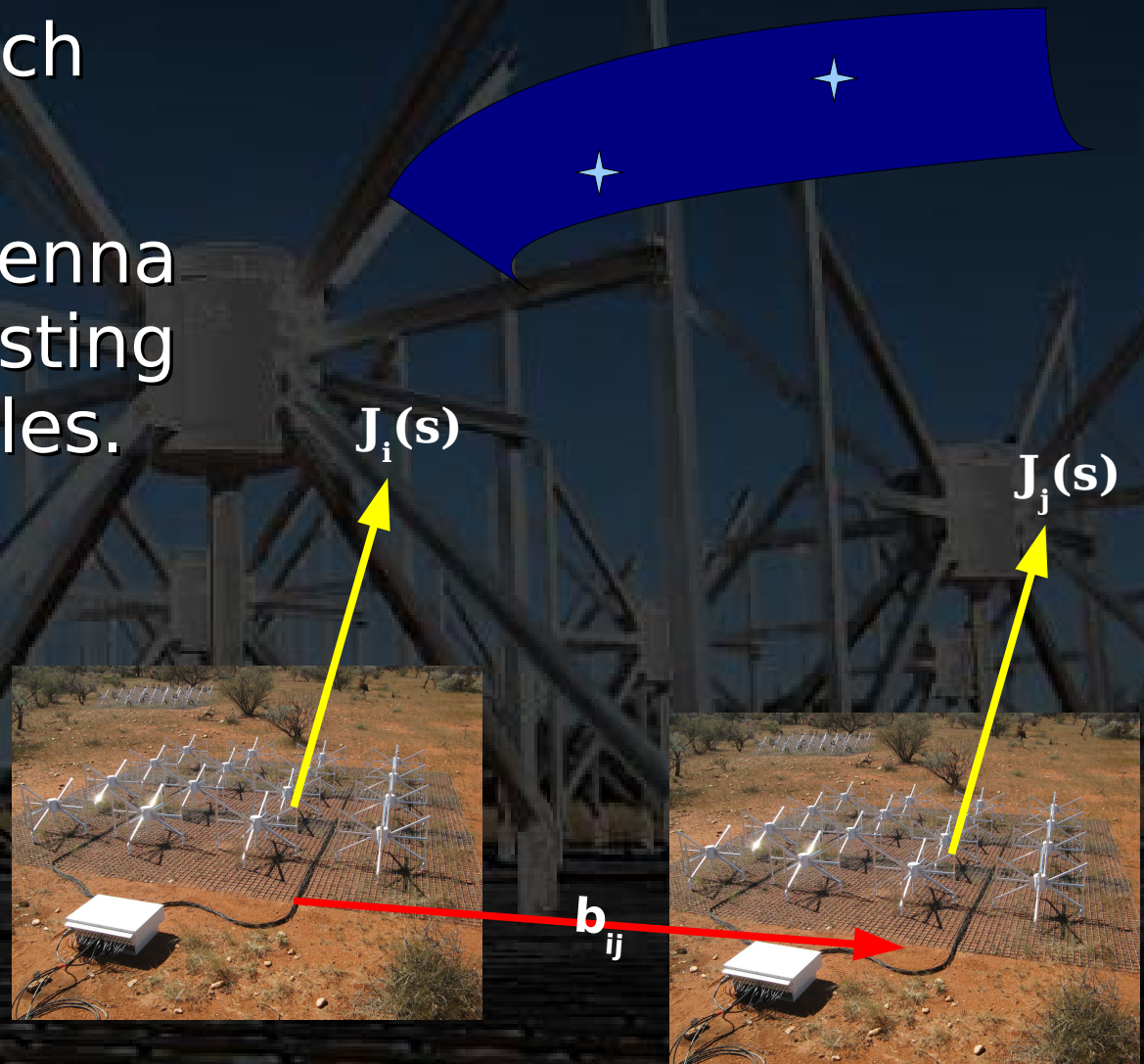


+RA onto
NS antenna

MAPS - point sources

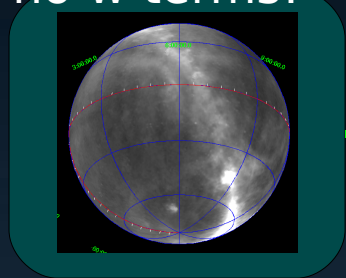
$$\mathbf{V}_{ij} = \mathbf{M}_{ij} \mathbf{M}_{ij}^S(\mathbf{s}) \mathbf{SI}(\mathbf{s}) e^{-2\pi i \mathbf{b}_{ij} \cdot (\mathbf{s} - \mathbf{s}_0)}$$

- Evaluate ME for each source...
- For MWA, each antenna is a “station” consisting of 16 dual-pol dipoles.



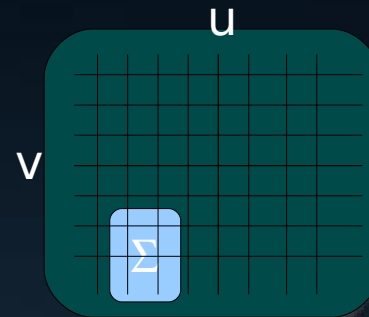
MAPS - diffuse sky

Entire sky centered on zenith
- no w terms!



Sky(l,m)
(I,Q,U,V)

FFT



Sky(u,v)
(I,Q,U,V)

Visibility (XX,XY,YX,YY)

Primary beam sample
(l,m)
32x32 pix

St 1(l,m)
Jones

St 2(l,m)
Jones

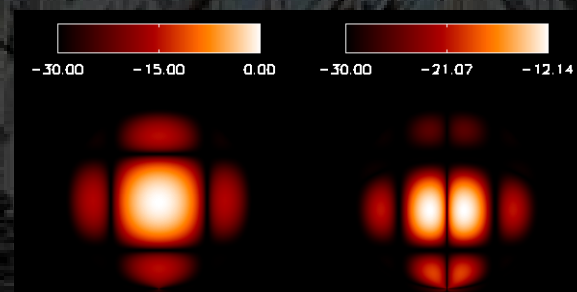


Baseline Mueller
(l,m)

...x16

FFT

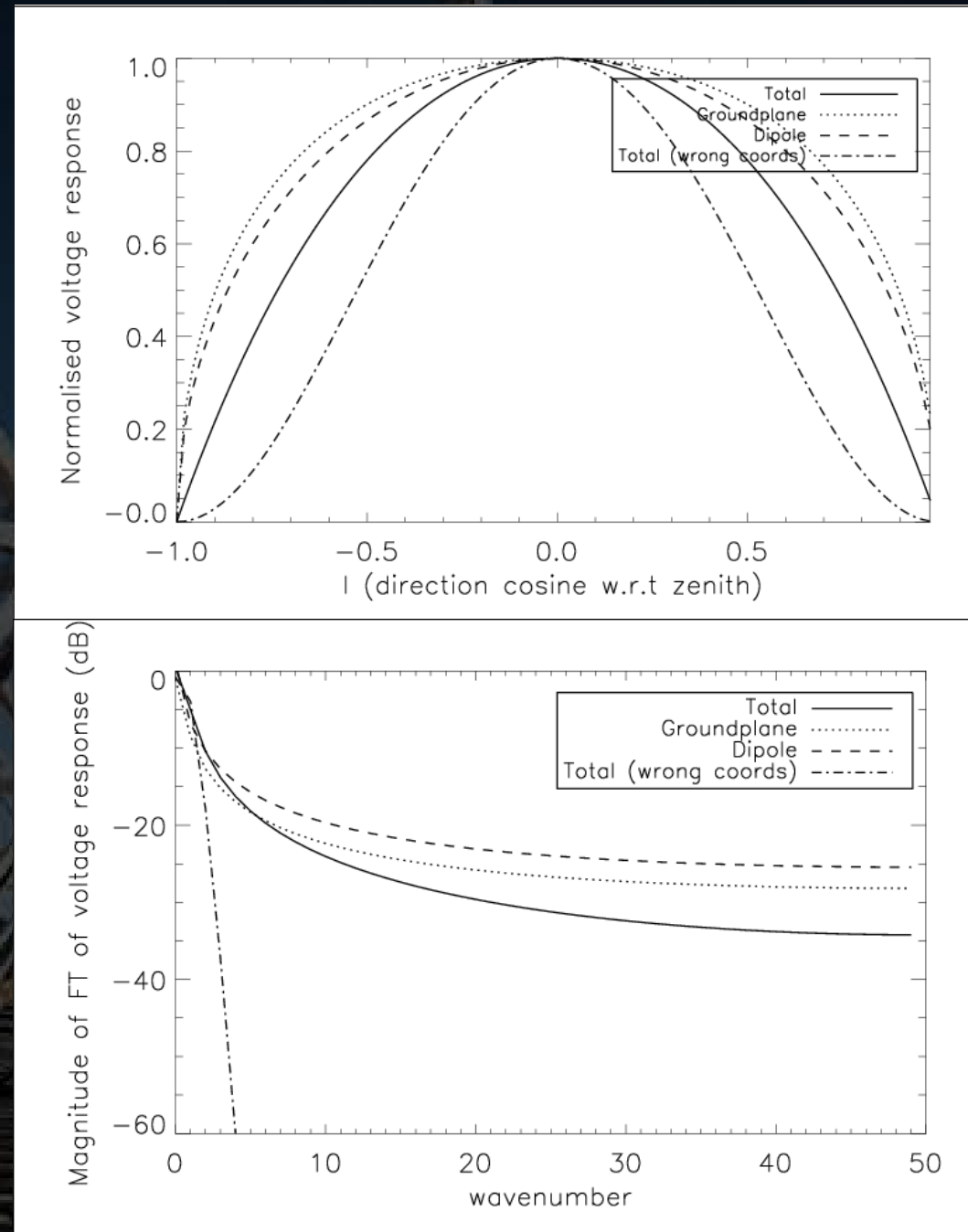
Baseline Mueller
(u,v)



- Time, freq averaging possible
- Ionosphere goes into station Jones matrix
- Sky sampling grid must adequately sample sky

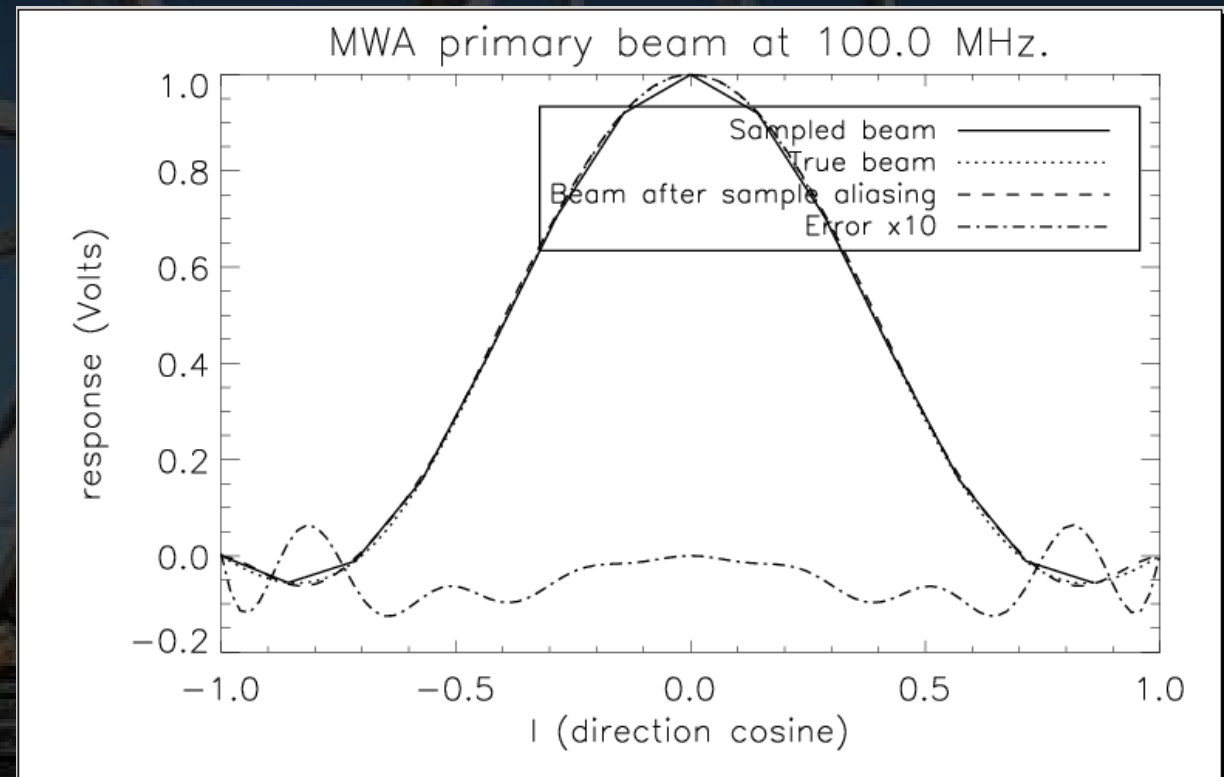
Primary beam convolution kernels (1)

- A short dipole over groundscreen does not have a (perfectly) compact (u,v) space representation
- This is a “wide field” effect: in the wrong coords (angle instead of direction cosine), the response is compact.



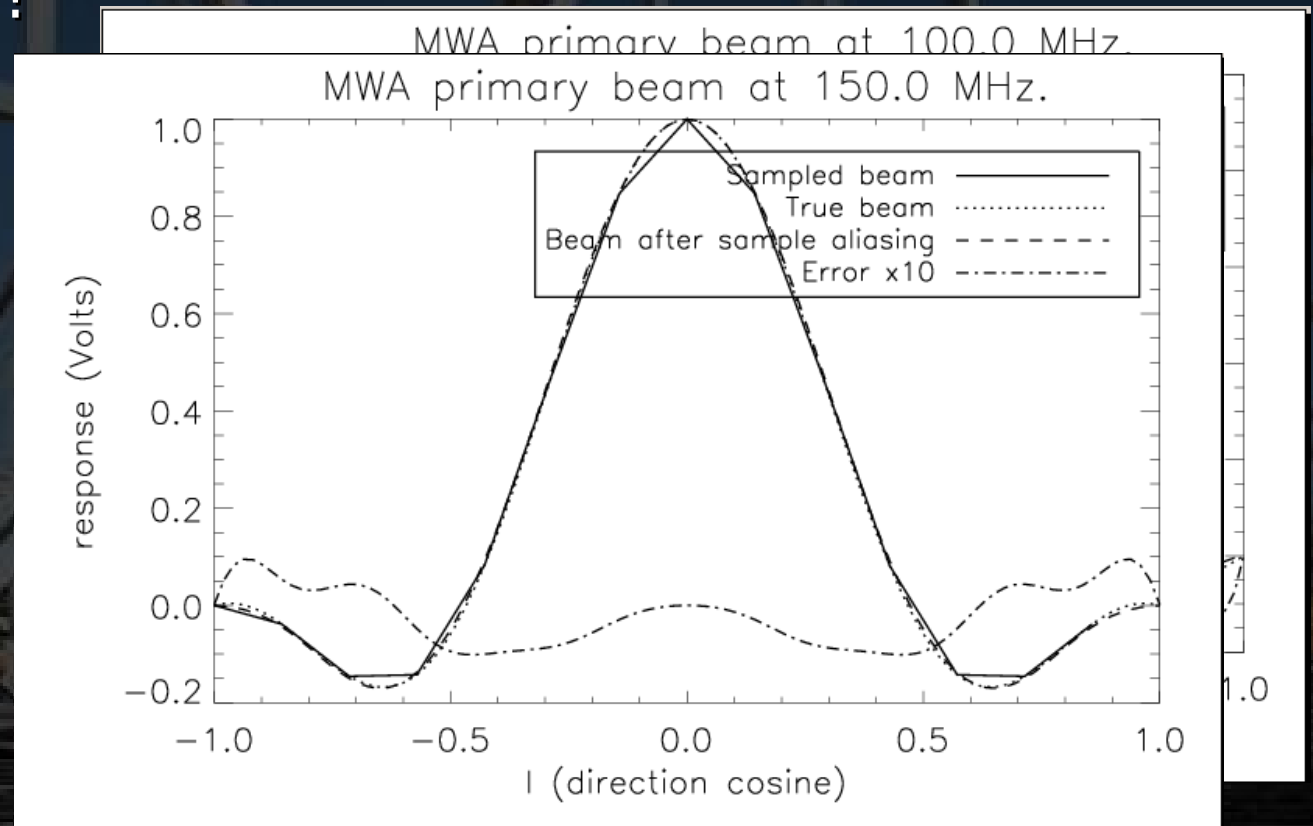
Primary beam convolution kernels (2)

- The MWA primary beam is non-negligible over the entire sky
- Sub-sampling the primary beam risks aliasing!



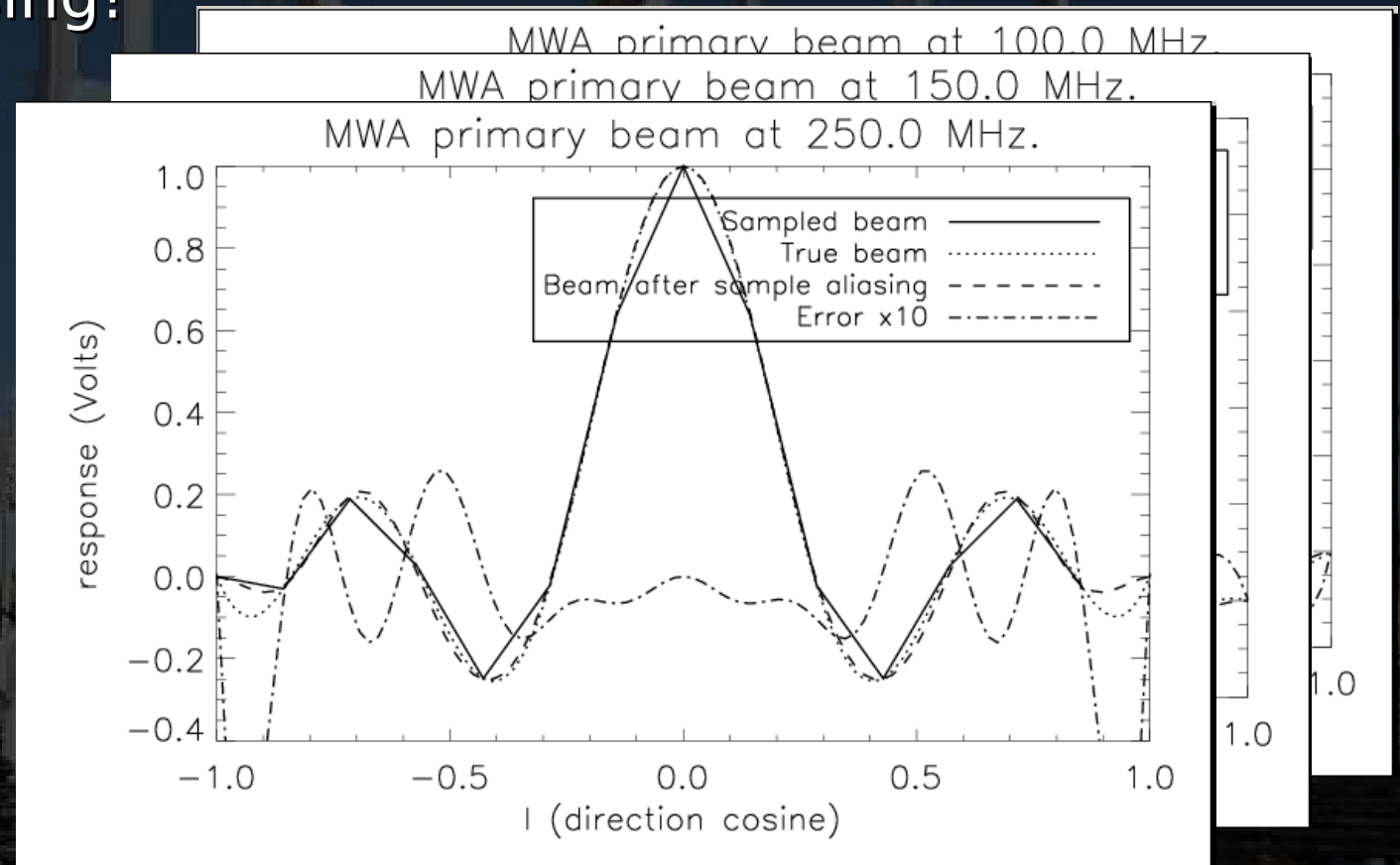
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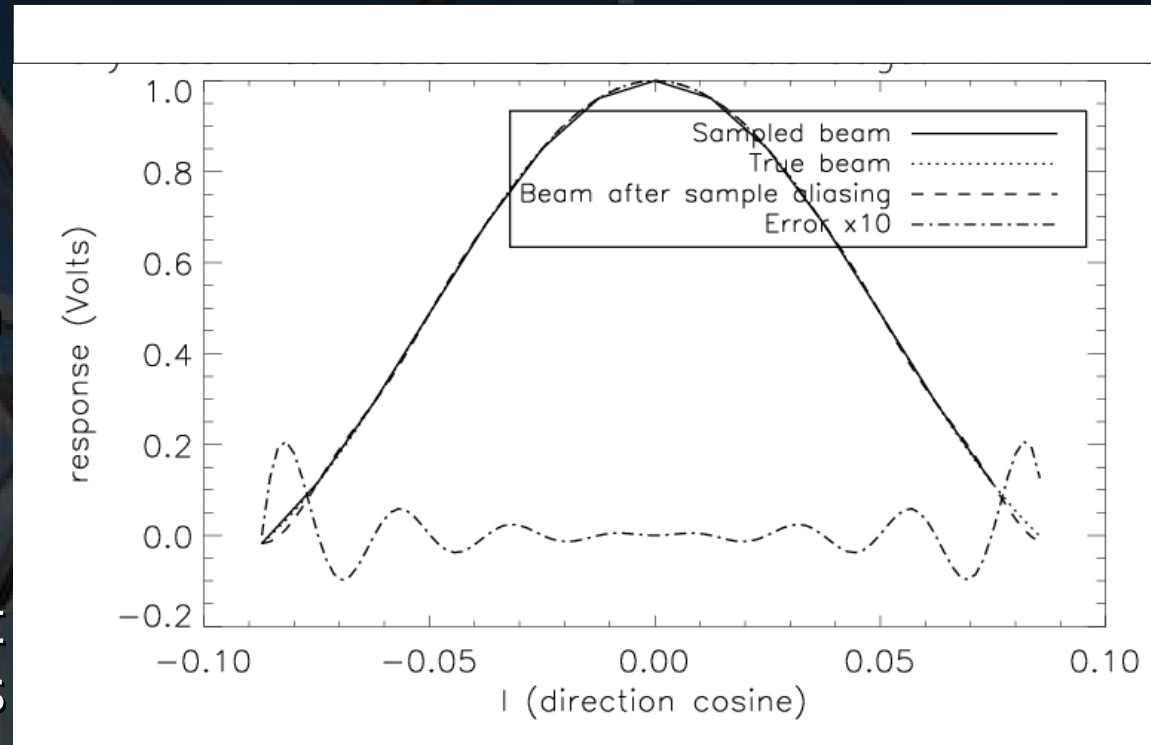
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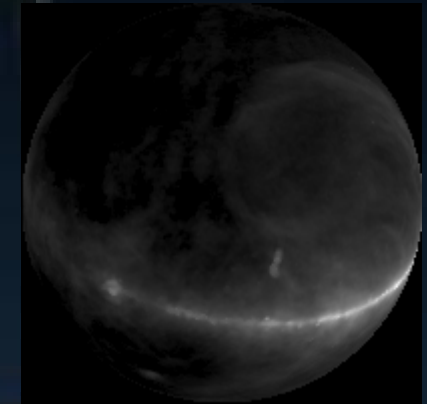
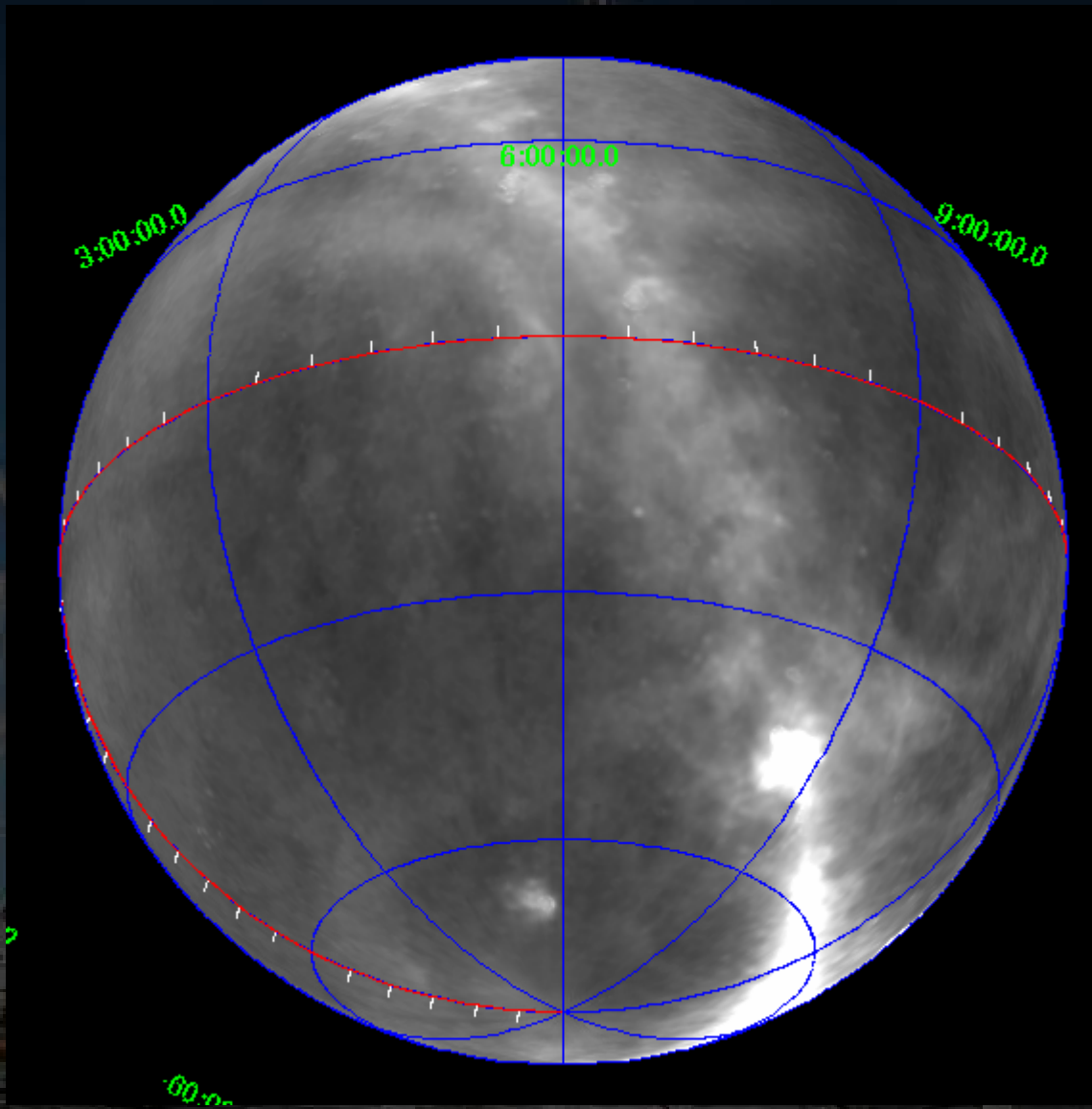


Primary beam convolution kernels (2)

- The MWA primary beam is non-negligible over the entire sky
- Sub-sampling the primary beam risks aliasing!
- Not just MWA... any time the beam is sampled with a discontinuous gradient will cause aliasing...
- E.g.: Uniformly illuminated circular aperture out to first null (FWHM = 4 degs in this case)



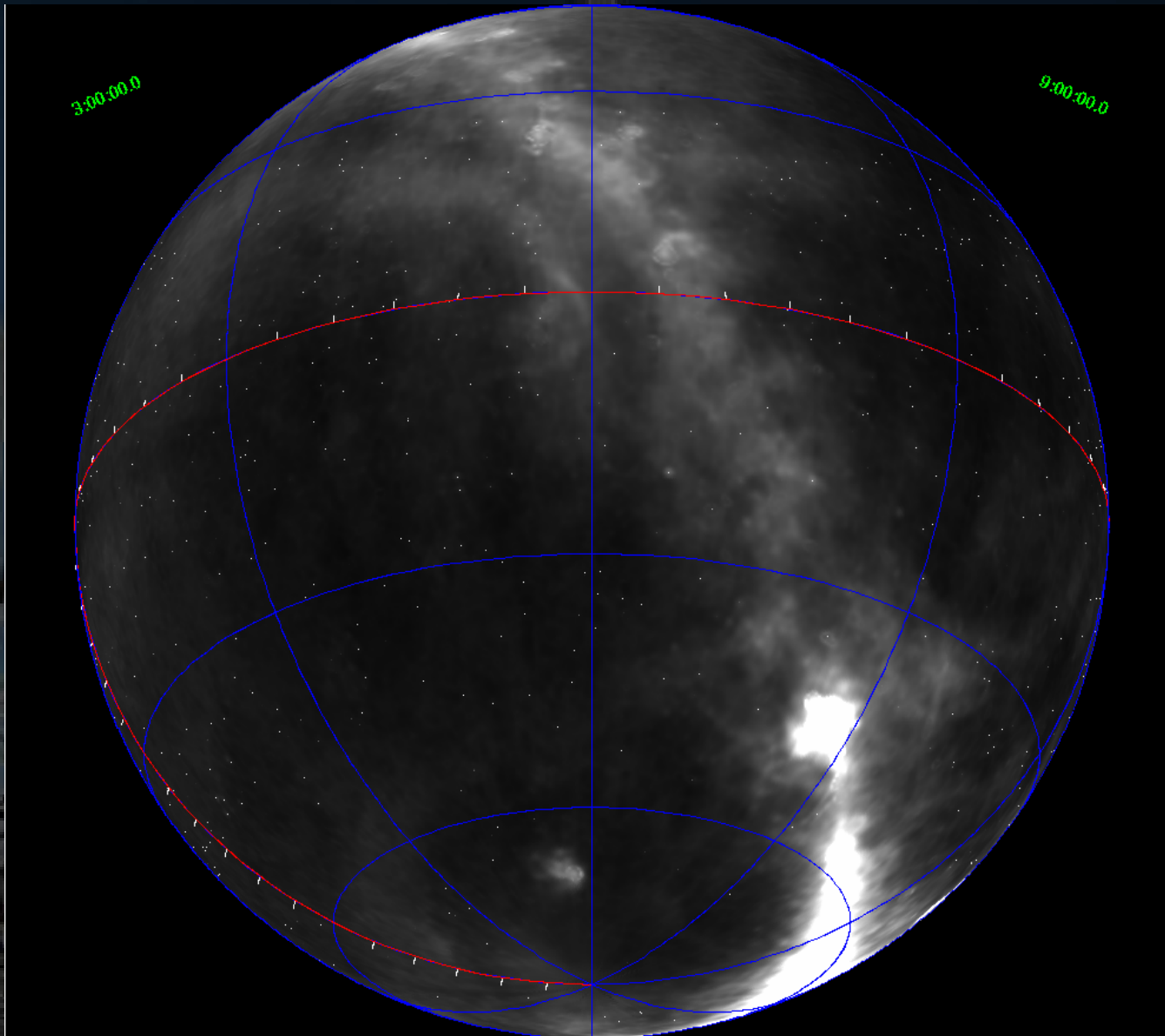
MAPS inputs: diffuse sky



Model 140MHz radio
sky at lat -26 deg

Diffuse sky model from Angelica de Oliveira-Costa
<http://space.mit.edu/home/angelica/gsm/>

Example - input sky for LST 6h

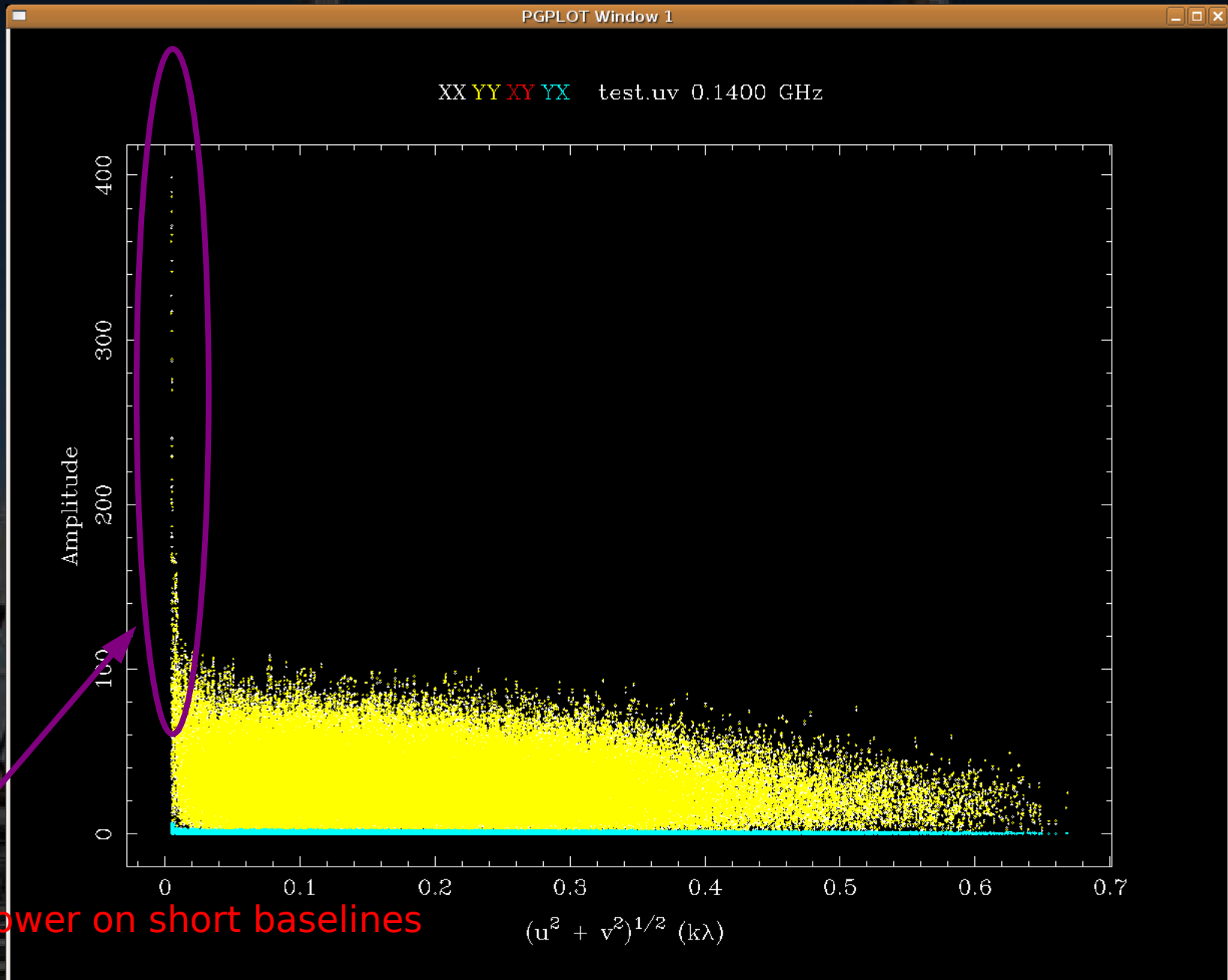


Point sources from
Parkes cat 90.

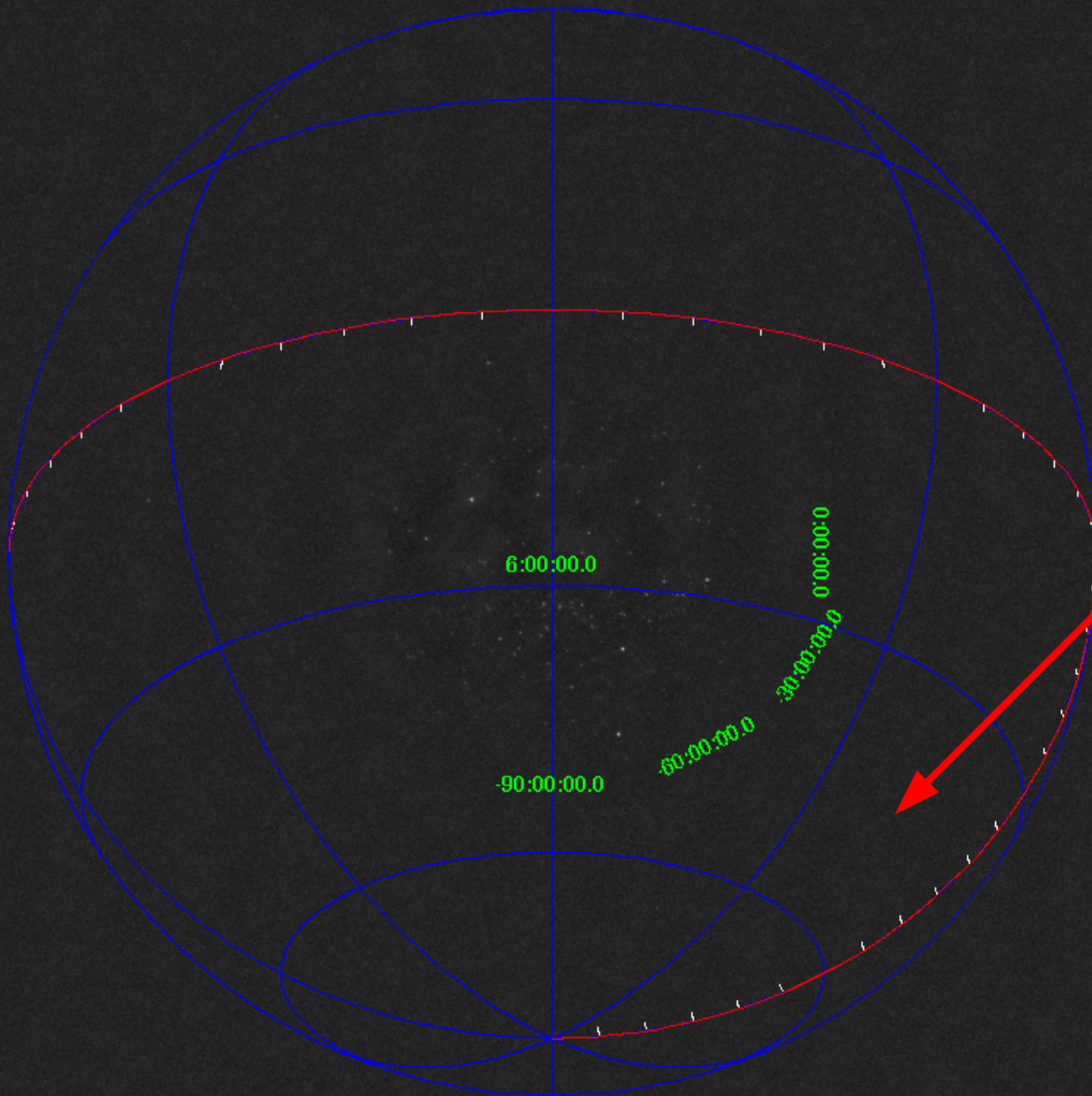
Pure Stokes I
input (in this
example)

Image size:
~6000x6000 pix

Example - 500T snapshot



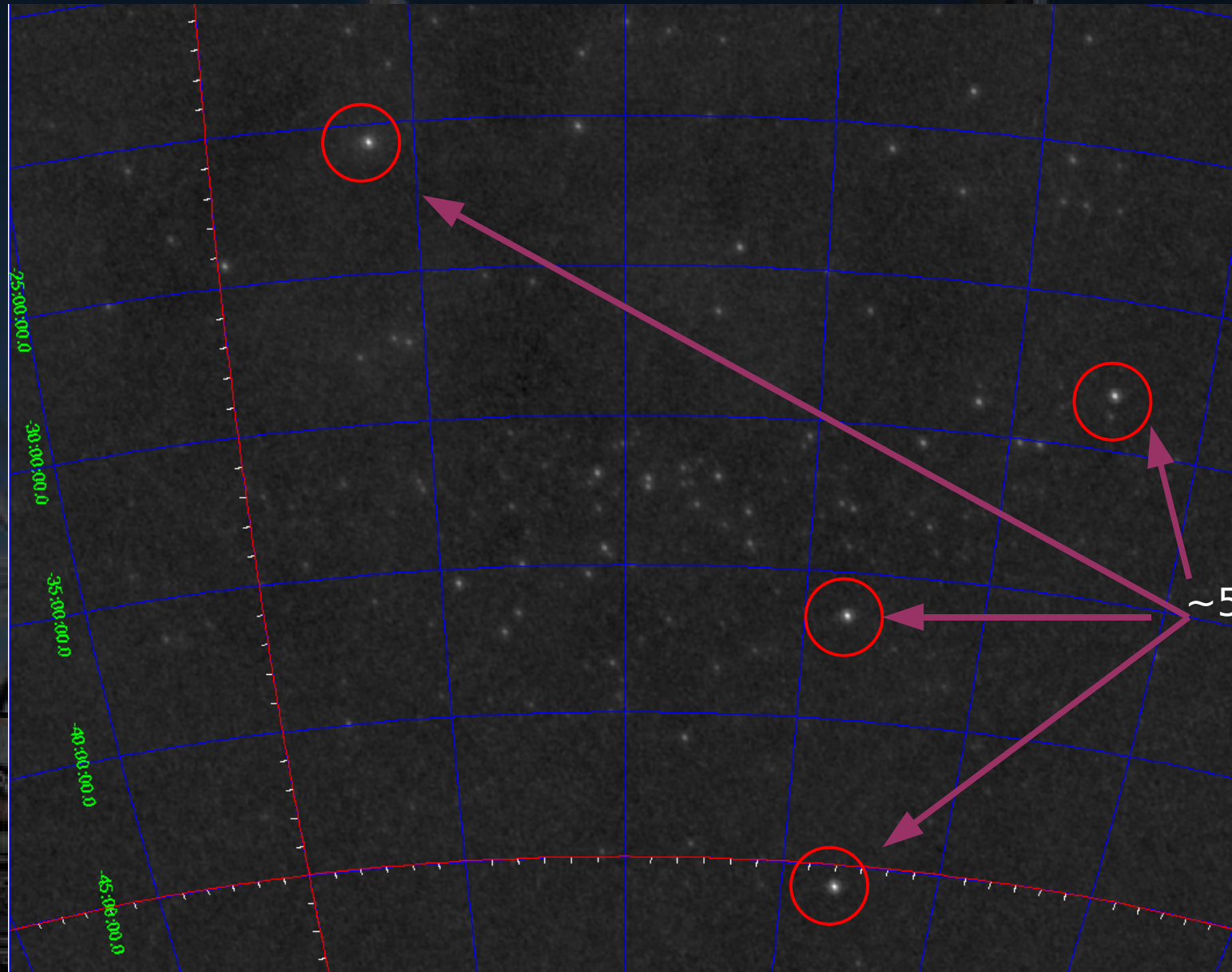
Example - 500T snapshot



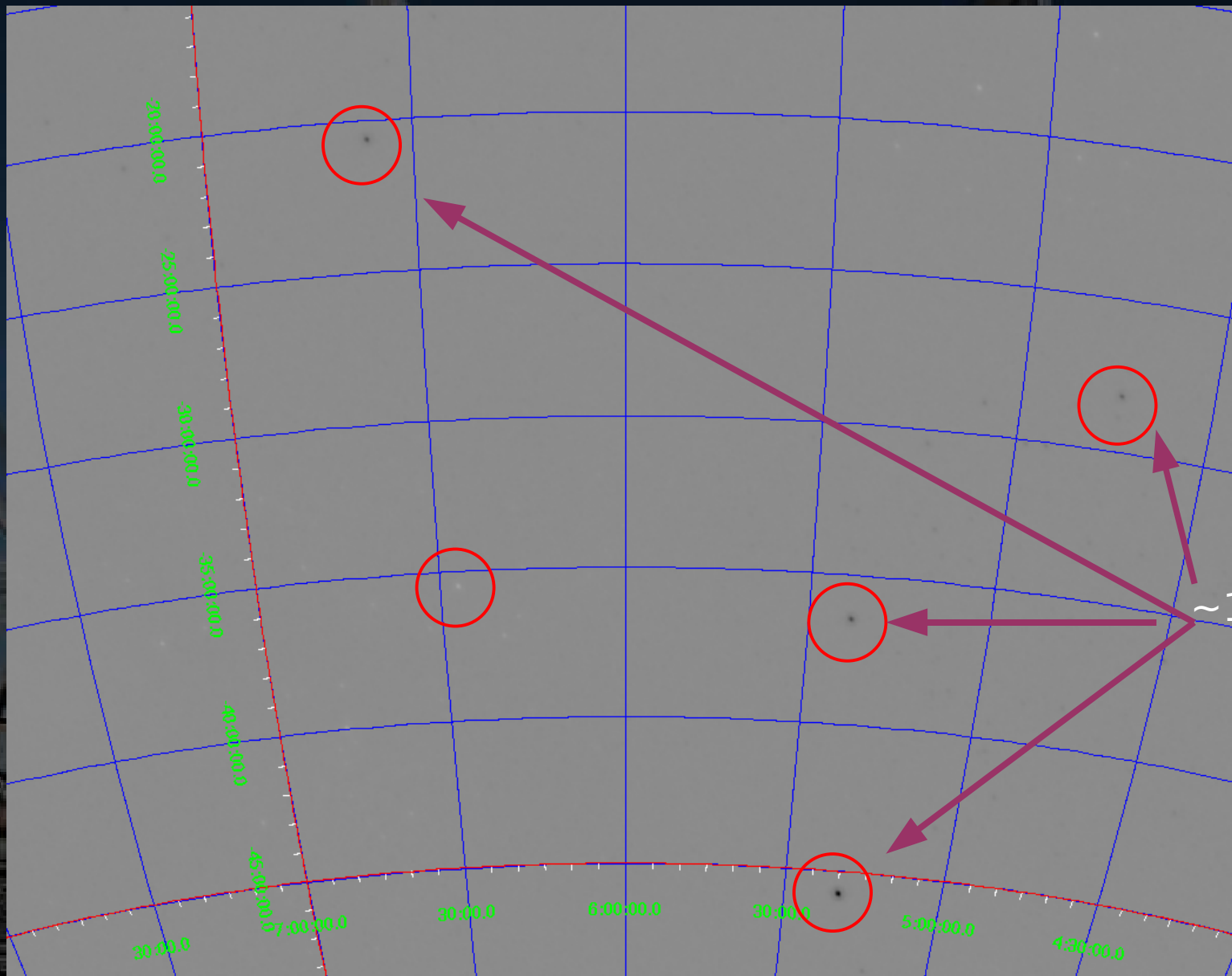
Note: sqrt stretch

Galactic plane is gone from antenna beams

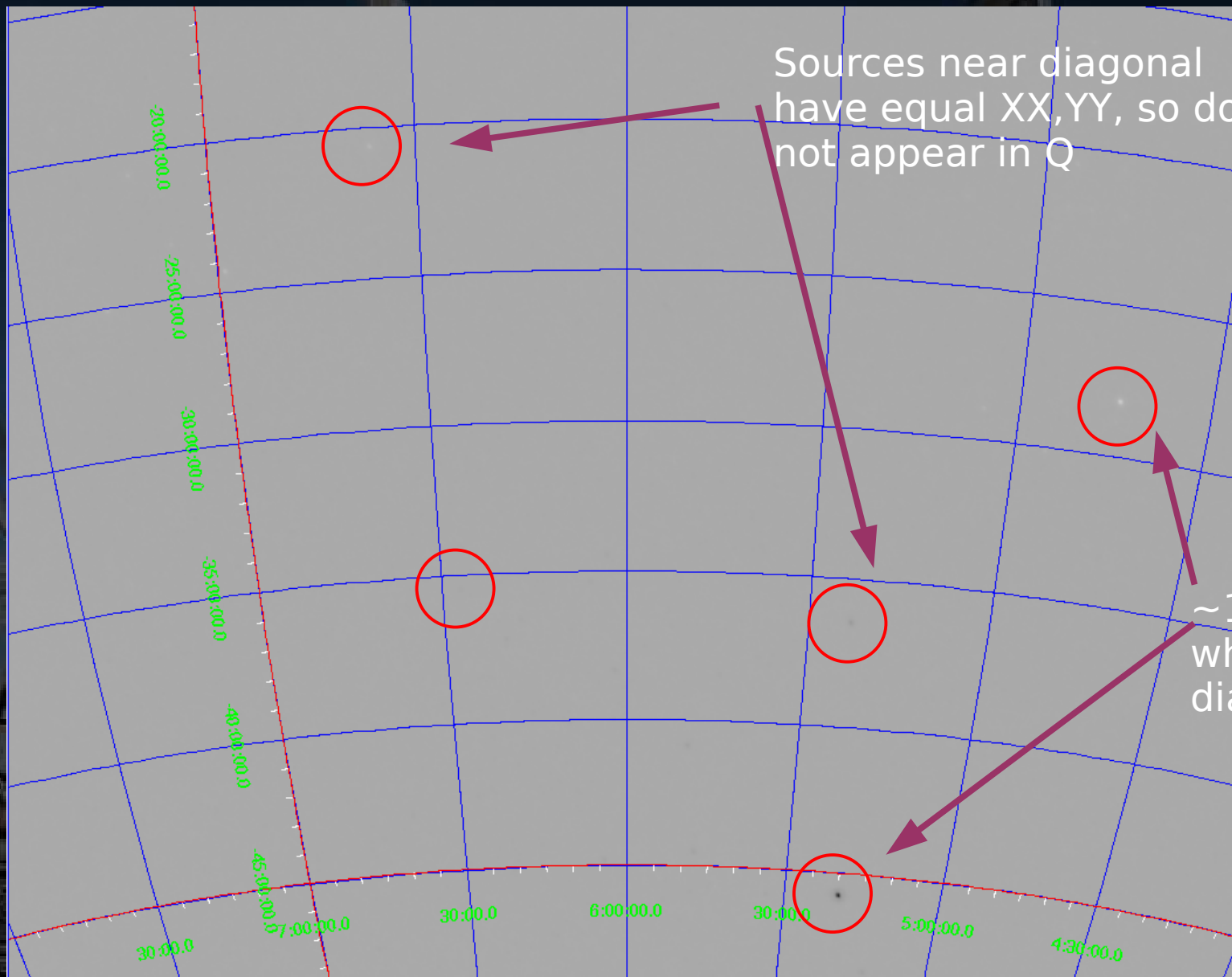
Example - 500T snapshot: Stokes I



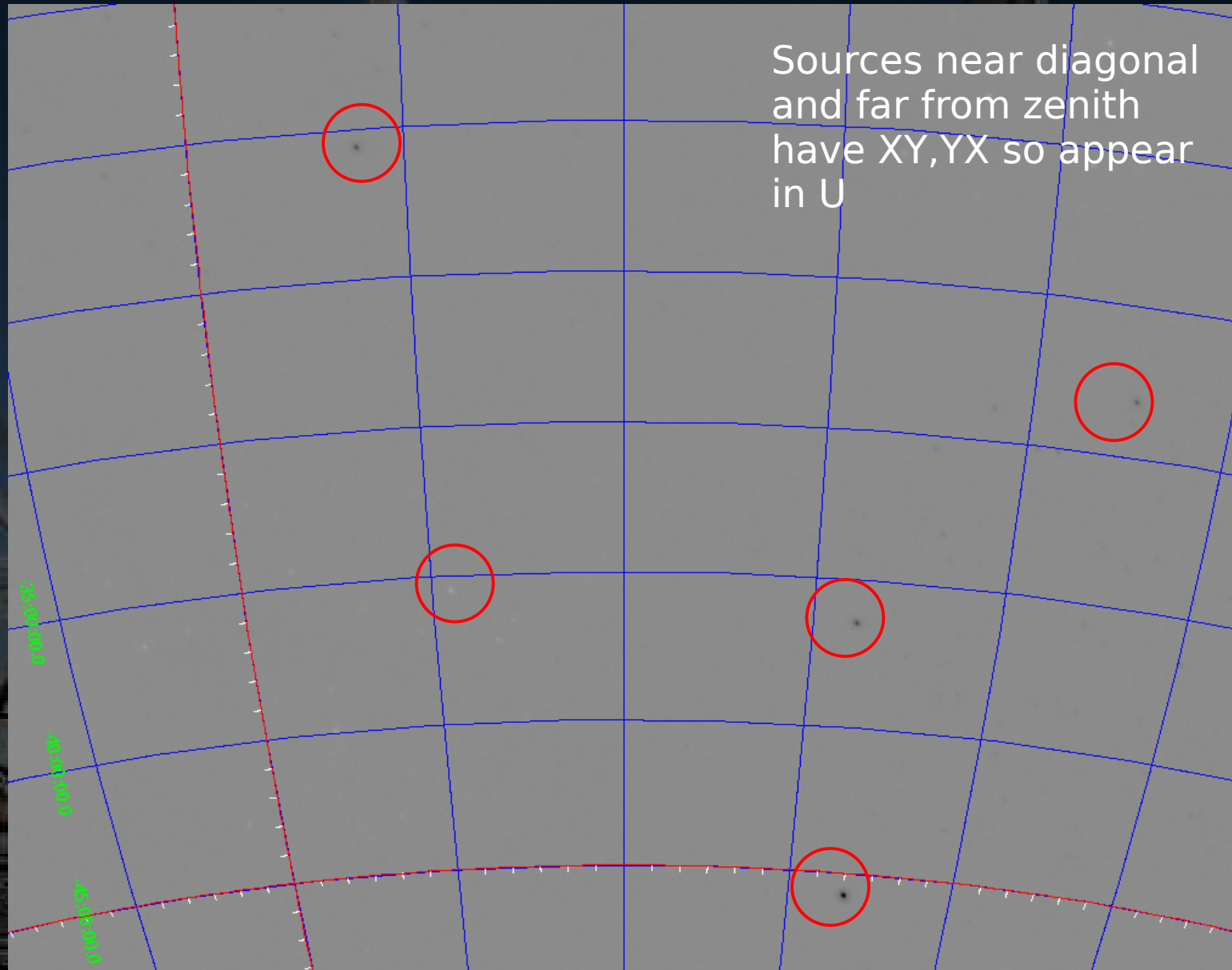
Example - 500T snapshot: XY



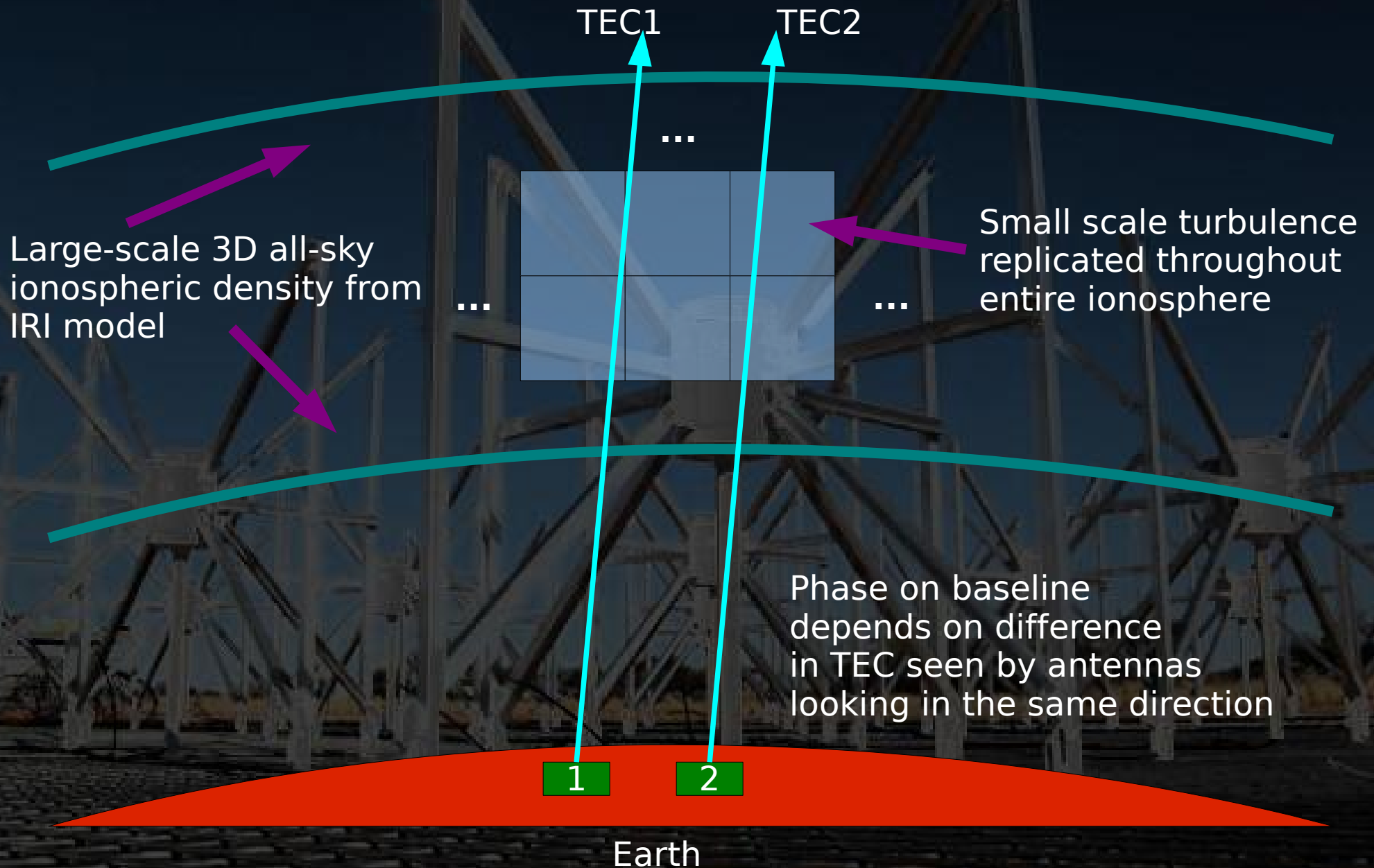
Example - 500T snapshot: Stokes 'Q'



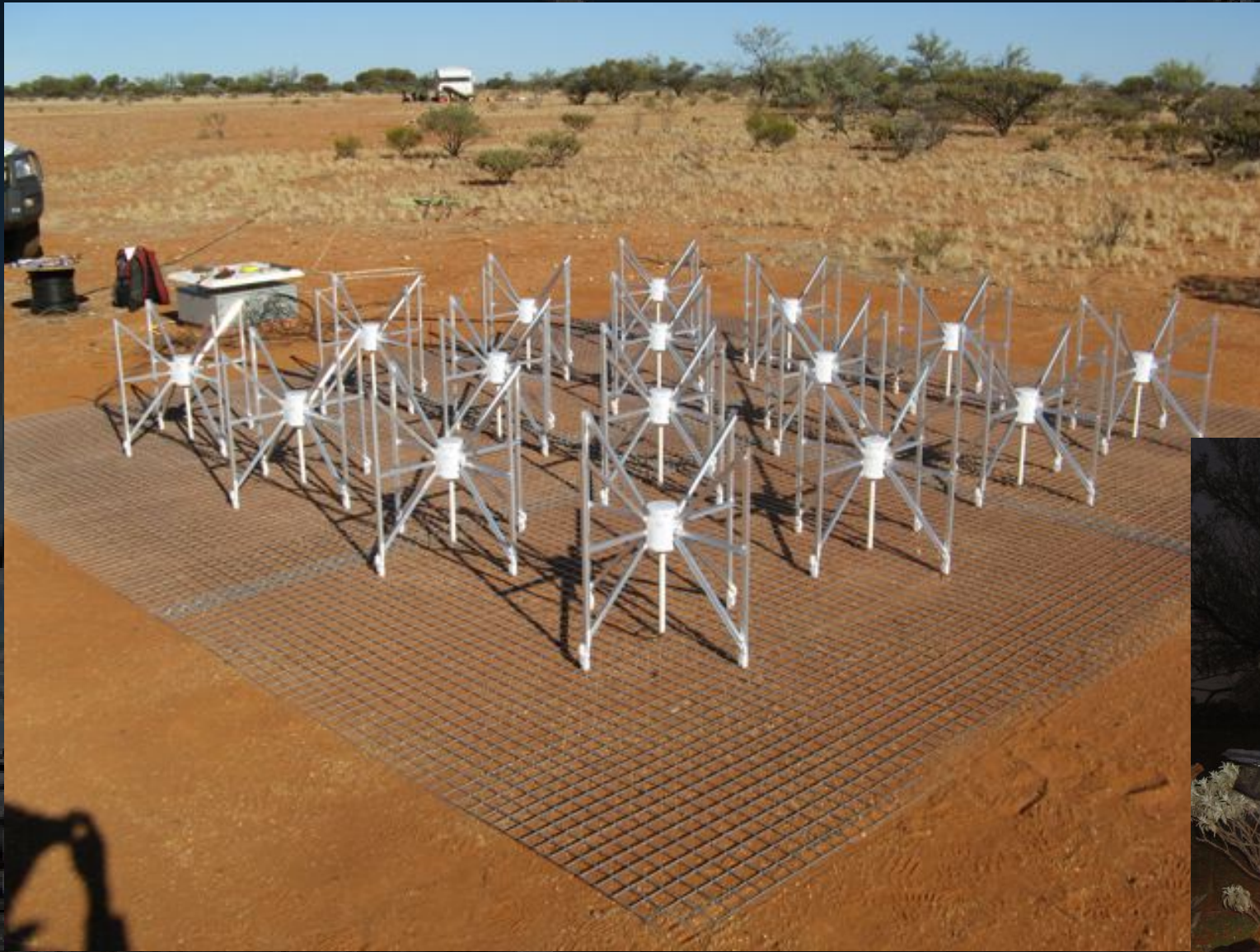
Example - 500T snapshot: Stokes 'U'



New and improved ionosphere



Questions?



10 Aug 2008

LF Software workshop